



# Carrier density noise in ET

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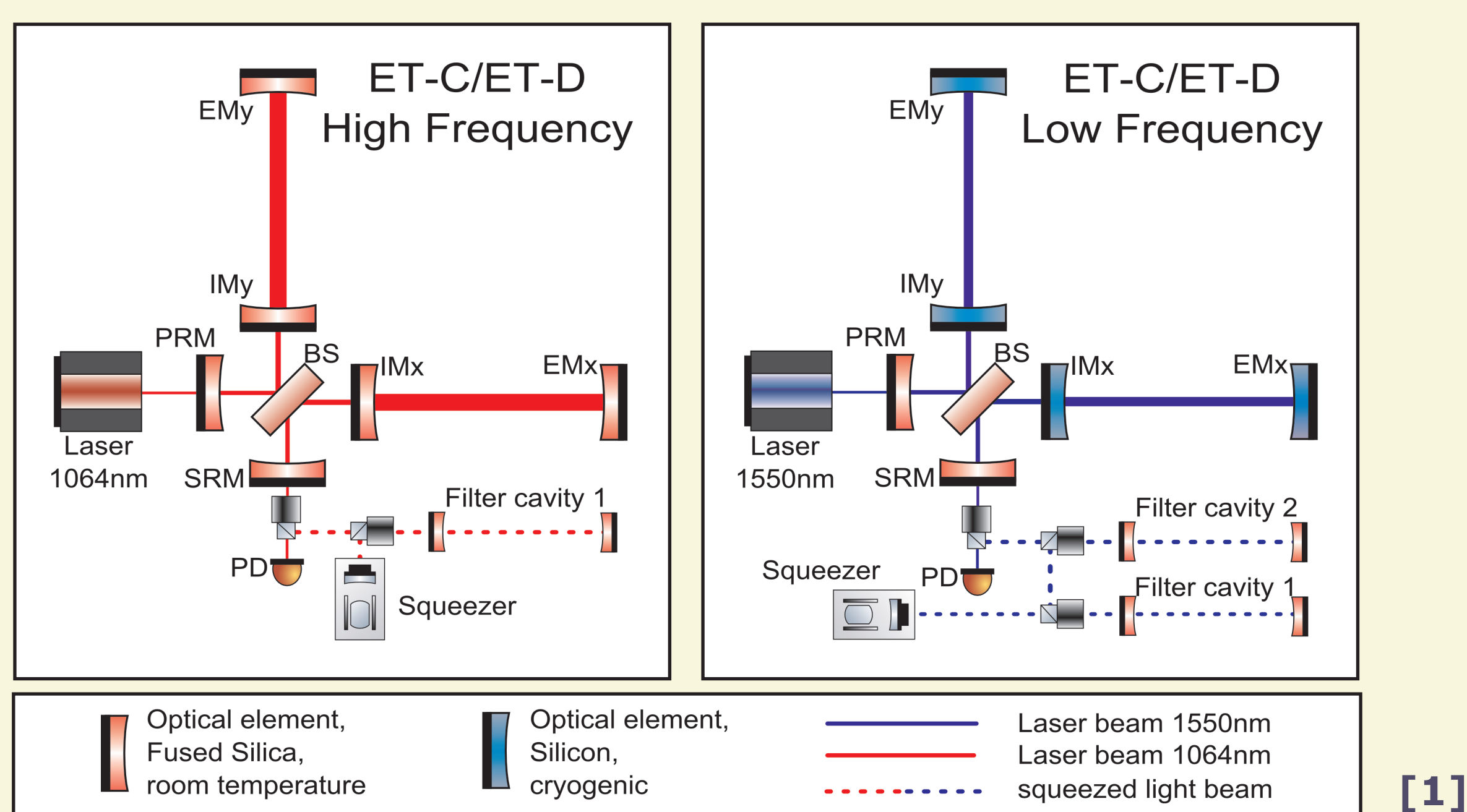
## Abstract

In contrast to fused silica as test mass material in current gravitational wave detectors, silicon represents a semiconductor. Thus, driven by thermal energy free electrons and holes are diffusing through the crystal. Via the effect of free carrier dispersion in silicon, i.e. the dependence of the refractive index on the free carrier density, such a diffusion introduces a noise process into transmissive optical components of the detector.

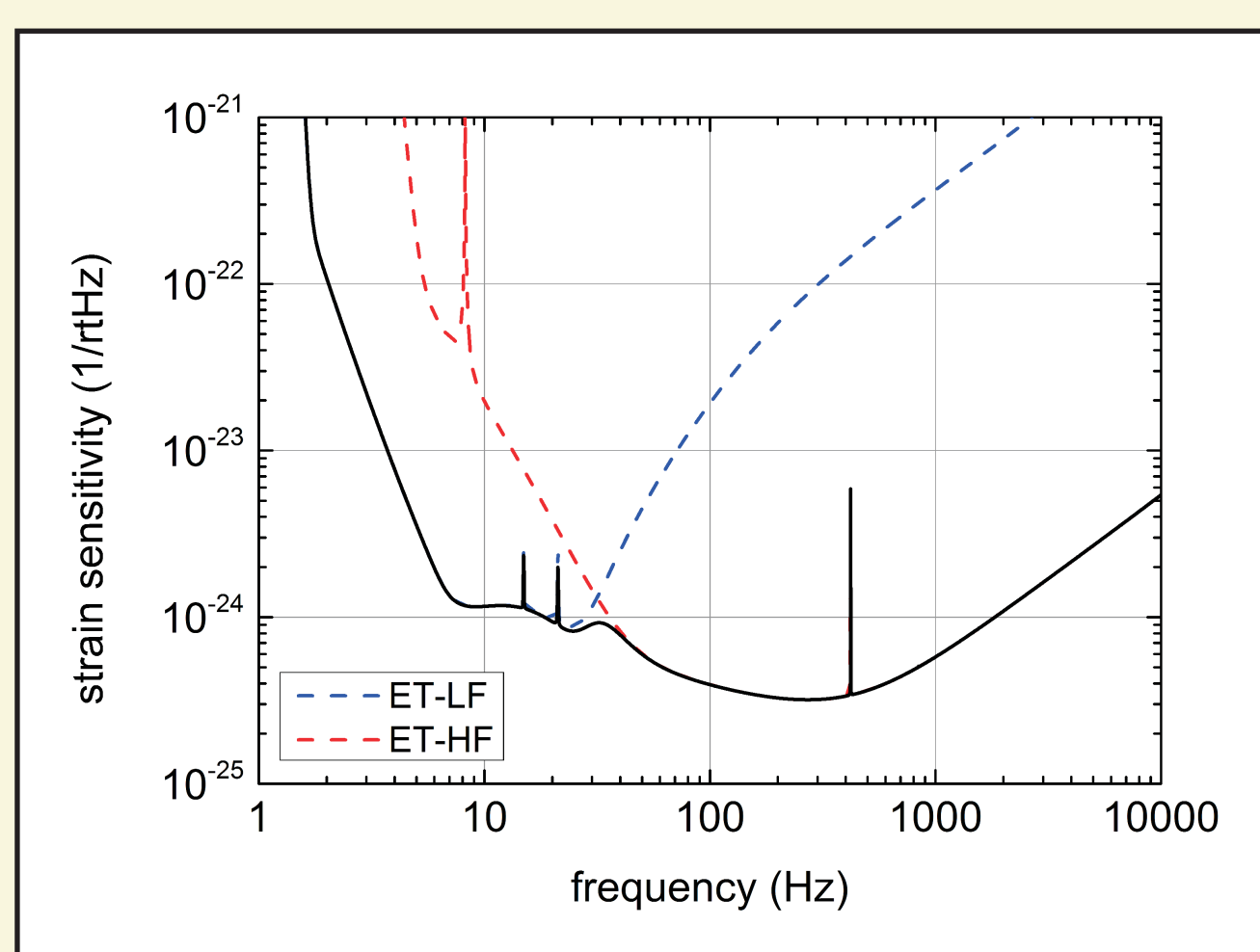
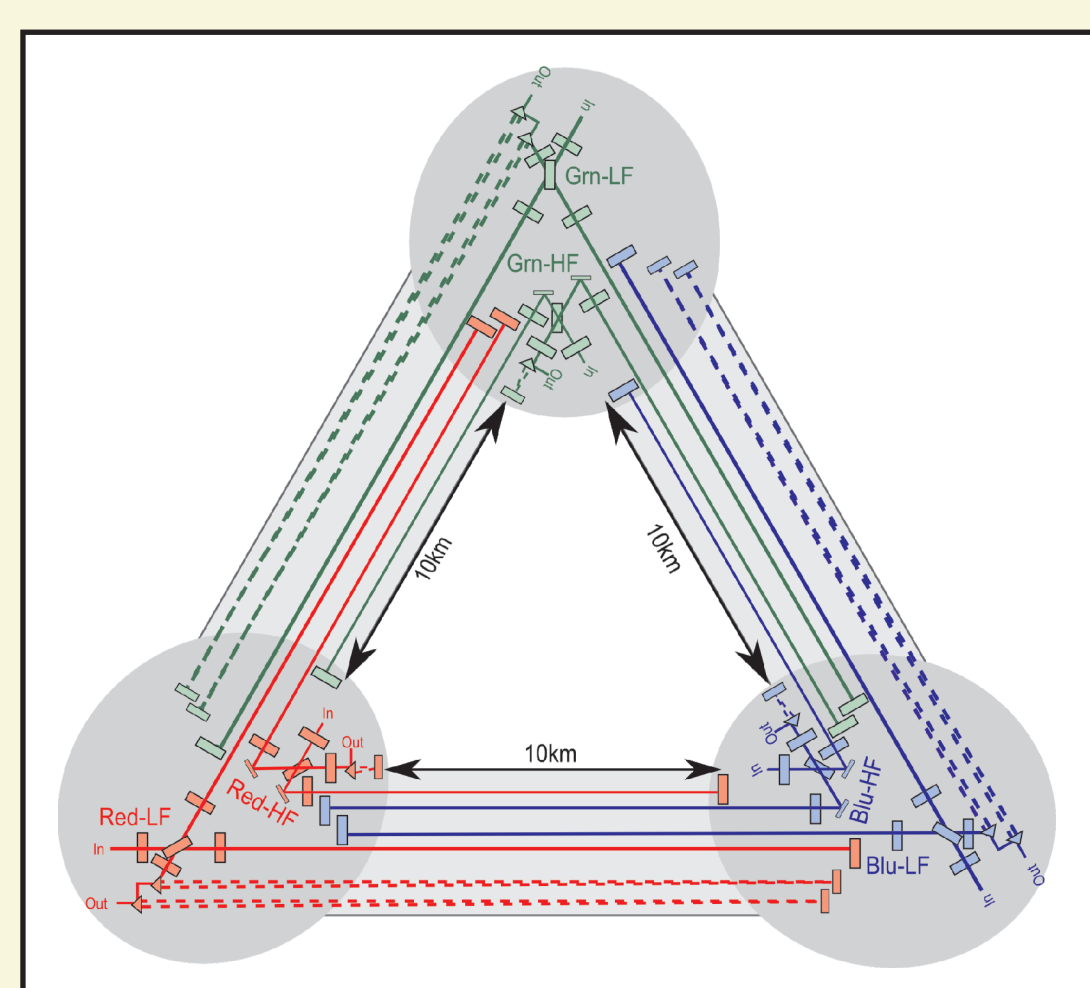
In this poster we present a theoretical analysis of carrier density noise and apply it to the proposed ET-LF design. Further, we evaluate its impact on the ET design goal and discuss possibilities to mitigate the effect of carrier density noise.

## Einstein Telescope

- Xylophone principle: ET-LF with cryogenic mirrors and moderate optical power  
ET-HF at room temperature and high optical power

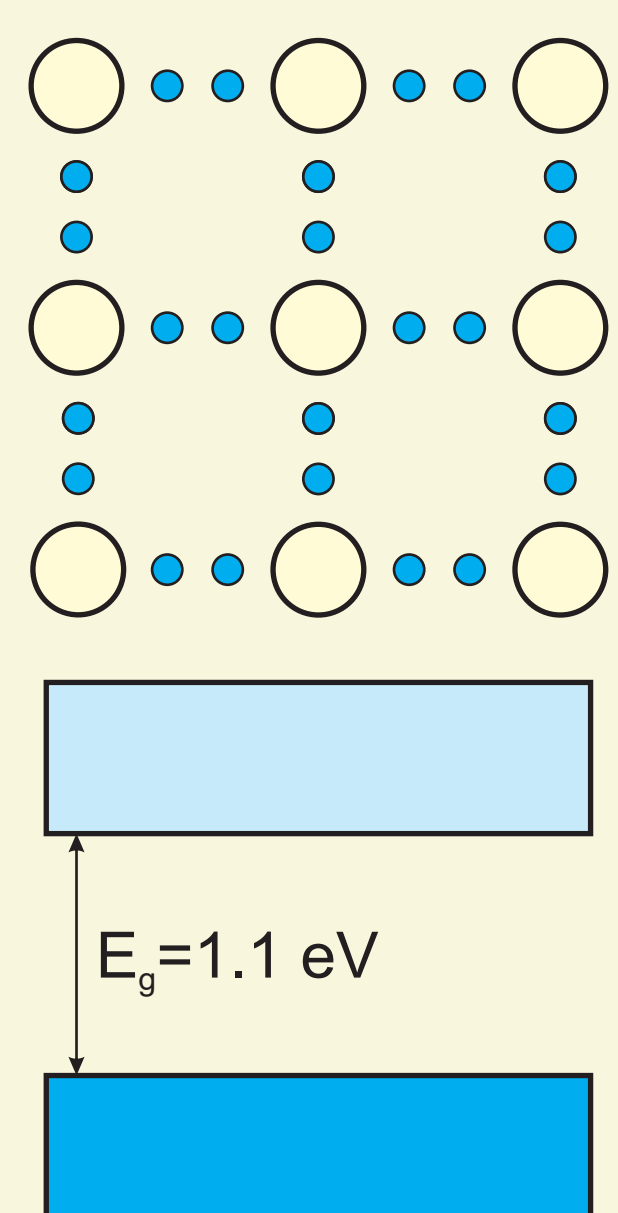


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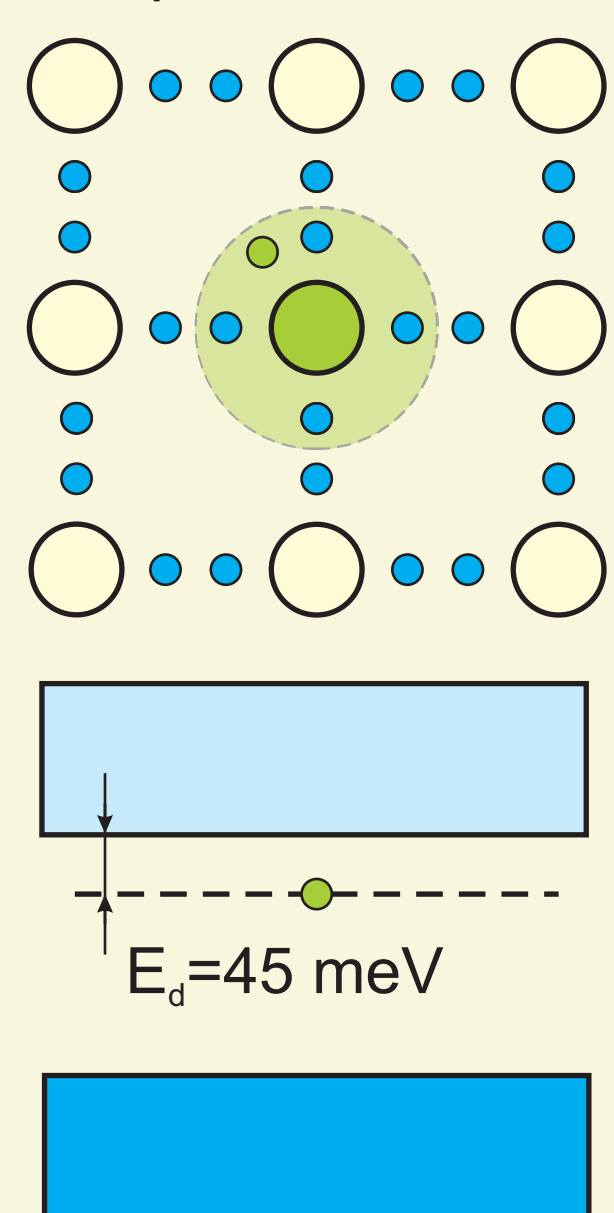


## Silicon as ET-LF substrate material

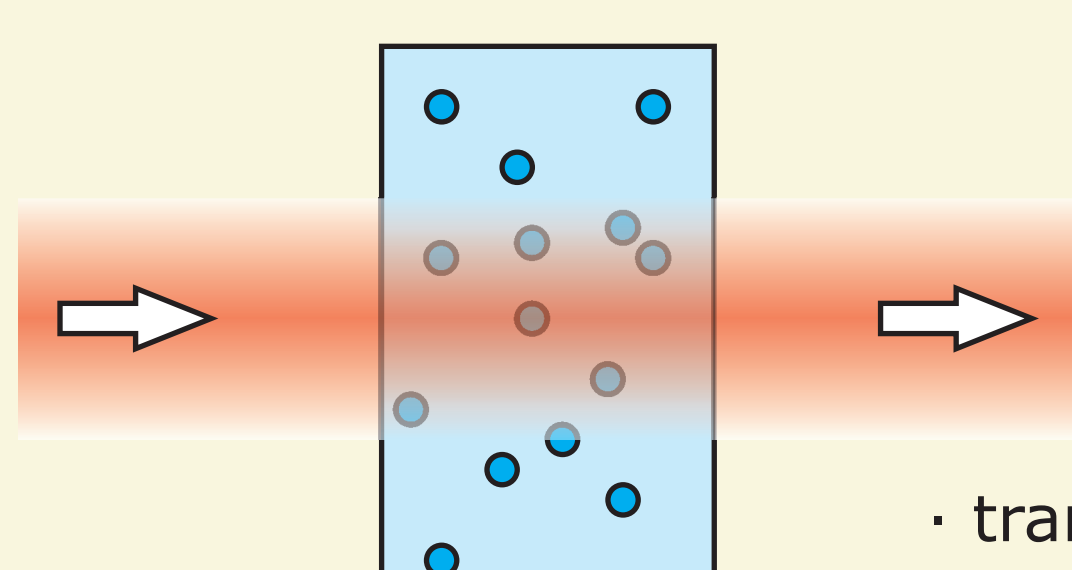
- intrinsic silicon



- n-doped silicon



- free carriers moving through the crystal



- transmitted light senses different number of free carriers

## Free Carrier Dispersion

- refractive index  $n$  is affected by the concentration of free holes  $n_h$  and electrons  $n_e$

$$n = -8.8 \times 10^{-22} \left( \frac{n_e}{\text{cm}^{-3}} \right) - 8.5 \times 10^{-18} \left( \frac{n_h}{\text{cm}^{-3}} \right)^{0.8} \quad [2]$$

## Diffusional CD noise

- Langevin approach for carrier concentration driven by random force  $F(r,t)$

$$\frac{\partial n}{\partial t}(r,t) - D \Delta n(r,t) = F(r,t)$$

- correlator of fluctuating forces

$$\langle F(\mathbf{k}, \omega) F^*(\mathbf{k}', \omega') \rangle = (2\pi)^4 F_0^2 k^2 \delta(\mathbf{k} - \mathbf{k}') \delta(\omega - \omega')$$

- normalization of  $F_0$  by thermodynamic equation

$$\langle \Delta N^2 \rangle_V = k_B T \left( \frac{\partial n}{\partial \mu} \right) V \quad [3]$$

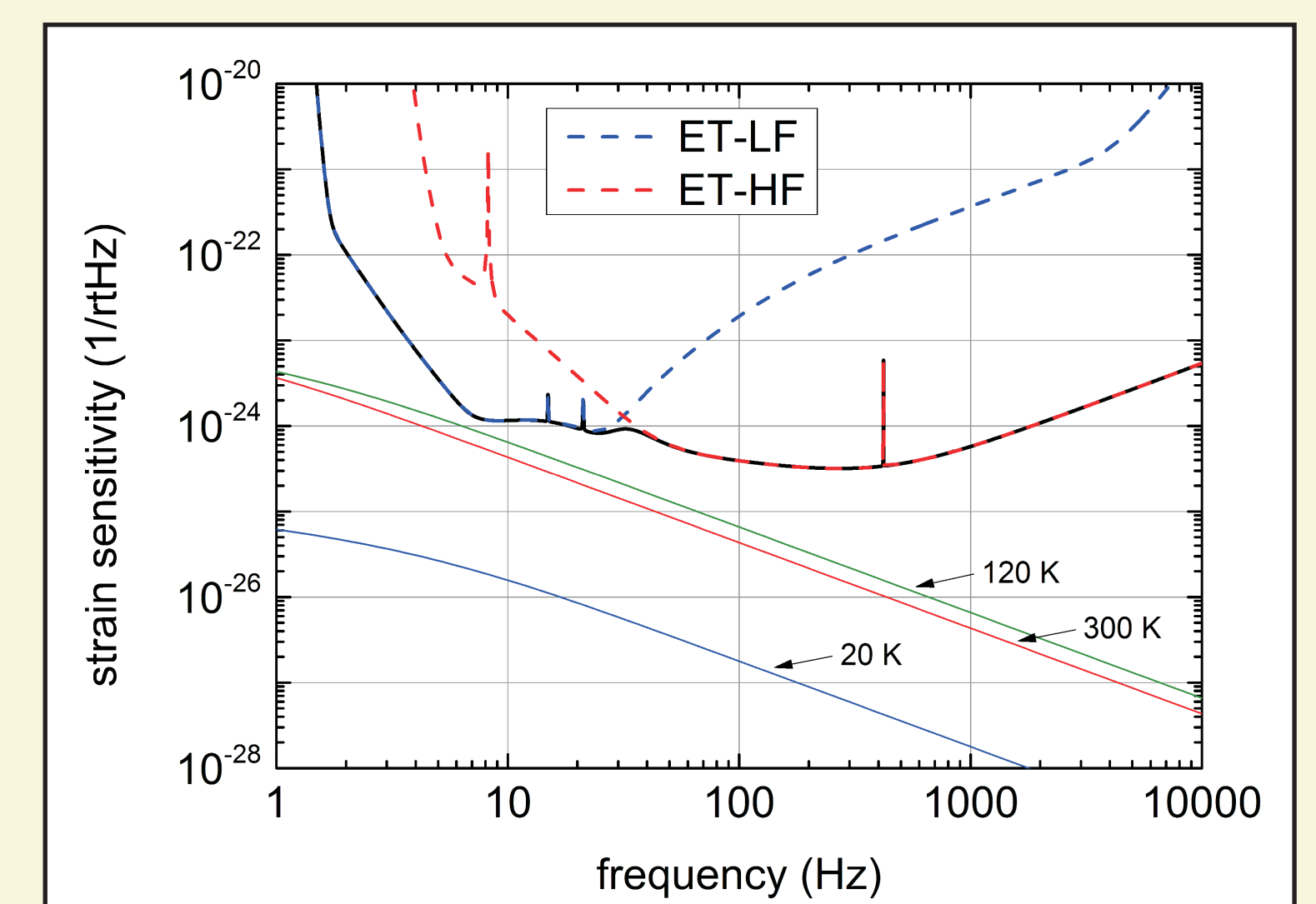
- displacement noise

$$S_z(\omega) = \frac{2}{\pi} k_B T \frac{\partial n}{\partial \mu} \gamma^2 H \int_0^\infty dk \frac{D k^3 e^{-k^2 r_0^2/2}}{\omega^2 + D^2 k^4}$$

- strain noise

$$S_h(\omega) \approx \frac{2S_z(\omega)}{L^2 \left( \frac{2F}{\pi} \right)^2}$$

- $H=46$  cm
- $w_0=9$  cm
- $L=10$  km
- $F=880$
- $n_{\text{dop}}=10^{14} \text{ cm}^{-3}$
- $\gamma_e=-8.8 \times 10^{-22} \text{ cm}^{-3}$



## Recombination CD noise

- even without diffusion ( $D=0$ ) excitation/recombination of free carriers induces CD noise

- Langevin approach using deviation from mean carrier density  $\eta$

$$\frac{\partial \eta}{\partial t}(r,t) + \frac{\eta(r,t)}{\tau_0} = F(r,t) \quad \text{with correlator}$$

$$\langle F(\mathbf{r}, t) F^*(\mathbf{r}', t') \rangle = F_0 \delta(\mathbf{r} - \mathbf{r}') \delta(t - t')$$

- normalization of  $F_0$  by

$$\langle \Delta N^2 \rangle_V = k_B T \left( \frac{\partial n}{\partial \mu} \right) V \approx nV$$

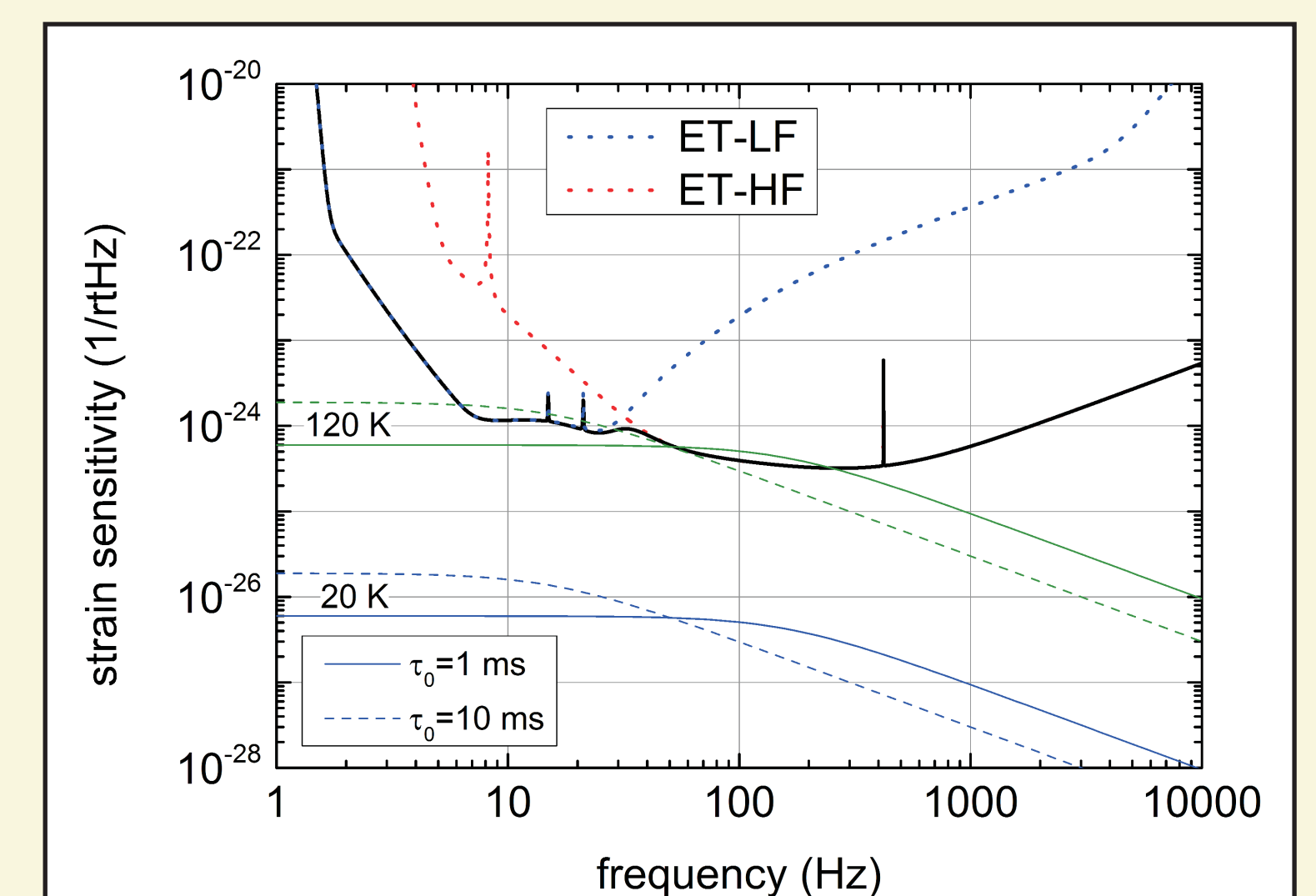
- displacement noise

$$S_z(\omega) = \frac{2H}{\pi r_0^2} \frac{\gamma^2 n \tau_0}{1 + \omega^2 \tau_0^2}$$

- recombination time  $\tau_0$

- $\tau_0 \sim 10$  ms for FZ Si
- $\tau_0 \sim 0.5$  ms for CZ Si

[4]



## References

- [1] ET science team, *Einstein gravitational wave Telescope conceptual design study*, ET-0106C-10, 2011, available at 'http://www.et-gw.eu/etdsdocument'
- [2] R. A. Soref and B. R. Bennett, *IEEE J. Quant. Electron QE-23*, 1987, pp. 123-129
- [3] L. D. Landau and E. M. Lifshitz, *Statistical Physics. Part 1*, Pergamon Press, Oxford, New York, 1980
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